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spectrum of that star had become nebular. Two lines at about the same wave-lengths were discovered in some of the well-known nebulæ by Dr. PALMER. In *Nova Geminorum* these lines are not yet developed, and are relatively very much weaker than they were in *Nova Persei*. As there are no early observations of this portion of the spectrum of *Nova Persei*, it is not possible to draw any conclusions as to these lines. It seems reasonable to expect, however, that they will develop along with the other nebular lines. C. D. PERRINE.

THE SPECTRUM OF *NOVA GEMINORUM*.

The spectrum of *Nova Geminorum* was observed, both visually and photographically at the Lick Observatory, with a spectroscope consisting of a single flint-prism mounted in conjunction with the collimator and camera of the regular Mills spectrograph. Unfortunately, unfavorable weather prevented any observations before April 1st, but after that date spectrograms were obtained on several evenings, both in the blue-violet region, and, by the use of isochromatic plates, in the visual part of the spectrum. The spectrum consists of broad, bright bands with some continuous spectrum. In its general features it resembles somewhat the spectrum of *Nova Persei* in March 1901, and also that in August of the following year; probably the star at its discovery had reached a stage in its development midway between those of *Nova Persei* at these two dates. Many of the bands coincide approximately with the lines of the hydrogen spectrum, and though a number of them could not be identified with hydrogen-lines, there was no certain trace of the so-called "nebular-lines," except the principal line λ 5007.

Between the date of the first spectrogram, April 2d, and that of the last, May 4th, the spectrum showed definite changes, but they are of such a nature that it cannot be said whether or not they indicate a progression toward the nebular type.

H. M. REESE.
H. D. CURTIS.

NOVA GEMINORUM.

The first opportunity of obtaining a long exposure on *Nova Geminorum* with the Crossley reflector, occurred on April 22d, 23d, and 24th, when a duration of 6^h 29^m was secured.

There are no indications of nebulosity about the *Nova* on this plate.

The spectrum was last photographed with the slitless spectrograph on May 11th. The star has faded nearly a magnitude in the interval, and the entire spectrum is weaker in consequence. There is a slight strengthening at about the position of the chief nebular line, and the bands at λ 373 and λ 350 seem to be slightly stronger. C. D. P.

May 21, 1903.

Two New Naked-eye Double Stars.

The brighter stars of the northern heavens have been examined so many times with telescopes of all sizes that one hardly expects to find among them a double star that has not been observed before. Two such stars, however, belonging to two very different classes of double stars, have apparently hitherto eluded detection. The first is a star in *Coma Berenices,* rated at 6.0 magnitude in the Bonn D. M. On April 8, 1903, I examined this star with the 12-inch telescope and found that it was attended by a 12th-magnitude companion. One night's measure with the 36-inch telescope gives the position-angle 259°.7 and the distance 1″.35. As the principal star has an annual proper motion of 0″.22 in 148°.5, both the angle and distance should increase if the two stars are not physically connected.

The other star is in the constellation *Boötis,* not far from *Arcturus,* and is given as 5.9 magnitude in the Harvard Photometry. With the 12-inch telescope, on May 5, 1903, this star was seen to be a very close double. Measures on two nights with the 36-inch give: position-angle, 199°.8; distance, 0″.20; magnitudes, 6.5–6.5. The star has no proper motion.

The positions of these two stars for 1900.0 are, respectively: R. A., $13^h\ 28^m\ 4^s$; Decl., $+24°\ 52'$; and R. A., $14^h\ 27^m\ 55^s$; Decl., $+27°\ 7'$. R. G. Aitken.

May 12, 1903.

Correction to the American Ephemeris for 1903.

In observing the satellites of *Uranus,* it has been my practice to guard against mistaking a small star for a satellite by computing the approximate positions of the satellites for